ESTIMATE OF SOLAR MAXIMUM USING THE 1–8 Å GEOSTATIONARY OPERATIONAL ENVIRONMENTAL SATELLITES X-RAY MEASUREMENTS
 (Postprint)

L. M. Winter and K. S. Balasubramaniam

12 December 2014

Journal Article

AIR FORCE RESEARCH LABORATORY
Space Vehicles Directorate
3550 Aberdeen Ave SE
AIR FORCE MATERIEL COMMAND
KIRTLAND AIR FORCE BASE, NM 87117-5776

REPORT DOCUMENTATION PAGE

1. REPORT DATE *(DD-MM-YYYY)*	2. REPORT TYPE	3. DATES COVERED *(From - To)*
12-12-2014	Journal Article	01 Oct 2014 – 30 Nov 2014

4. TITLE AND SUBTITLE	5a. CONTRACT NUMBER
Estimate of Solar Maximum Using the 1–8 Å *Geostationary Operational Environmental Satellites* X-Ray Measurements (Postprint)	
	5b. GRANT NUMBER
	5c. PROGRAM ELEMENT NUMBER 61102F

6. AUTHOR(S)	5d. PROJECT NUMBER
L. M Winter and K. S. Balasubramaniam	3001
	5e. TASK NUMBER PPM00016108
	5f. WORK UNIT NUMBER EF122003

7. PERFORMING ORGANIZATION NAME(S) AND ADDRESS(ES)	8. PERFORMING ORGANIZATION REPORT NUMBER
Air Force Research Laboratory Space Vehicles Directorate 3550 Aberdeen Avenue SE Kirtland AFB, NM 87117-5776	AFRL-RV-PS-TR-2015-0005

9. SPONSORING / MONITORING AGENCY NAME(S) AND ADDRESS(ES)	10. SPONSOR/MONITOR'S ACRONYM(S) AFRL/RVBXS
	11. SPONSOR/MONITOR'S REPORT NUMBER(S)

12. DISTRIBUTION / AVAILABILITY STATEMENT
Approved for public release; distribution is unlimited. (377ABW-2014-0608 dtd 01 Aug 2014)

13. SUPPLEMENTARY NOTES

14. ABSTRACT
We present an alternate method of determining the progression of the solar cycle through an analysis of the solar X-ray background. Our results are based on the NOAA *Geostationary Operational Environmental Satellites* (*GOES*) X-ray data in the 1–8 Å band from 1986 to the present, covering solar cycles 22, 23, and 24. The X-ray background level tracks the progression of the solar cycle through its maximum and minimum. Using the X-ray data, we can therefore make estimates of the solar cycle progression and the date of solar maximum. Based upon our analysis, we conclude that the Sun reached its hemisphere-averaged maximum in solar cycle 24 in late 2013. This is within six months of the NOAA prediction of a maximum in spring 2013.

15. SUBJECT TERMS
Solar flares, x-ray flux, flare, solar cycle

16. SECURITY CLASSIFICATION OF:			17. LIMITATION OF ABSTRACT	18. NUMBER OF PAGES	19a. NAME OF RESPONSIBLE PERSON Dr. K. S. Balasubramaniam
a. REPORT Unclassified	b. ABSTRACT Unclassified	c. THIS PAGE Unclassified	Unlimited	8	19b. TELEPHONE NUMBER *(include area code)*

ESTIMATE OF SOLAR MAXIMUM USING THE 1–8 Å *GEOSTATIONARY OPERATIONAL ENVIRONMENTAL SATELLITES* X-RAY MEASUREMENTS

L. M. WINTER[1] AND K. S. BALASUBRAMANIAM[2]
[1] Atmospheric and Environmental Research, 131 Hartwell Avenue, Lexington, MA 02421, USA; lwinter@aer.com
[2] Space Weather Effects Section, Space Vehicles Directorate, Air Force Research Laboratory, Kirtland AFB, NM 87117, USA

ABSTRACT

We present an alternate method of determining the progression of the solar cycle through an analysis of the solar X-ray background. Our results are based on the NOAA *Geostationary Operational Environmental Satellites* (*GOES*) X-ray data in the 1–8 Å band from 1986 to the present, covering solar cycles 22, 23, and 24. The X-ray background level tracks the progression of the solar cycle through its maximum and minimum. Using the X-ray data, we can therefore make estimates of the solar cycle progression and the date of solar maximum. Based upon our analysis, we conclude that the Sun reached its hemisphere-averaged maximum in solar cycle 24 in late 2013. This is within six months of the NOAA prediction of a maximum in spring 2013.

Key words: Sun: activity – Sun: X-rays, gamma rays

Online-only material: color figures

1. INTRODUCTION

Predictions of the length of the solar cycle and the date of solar maximum are important for planning space missions and satellite orbits. Besides direct electromagnetic, particle, and mass effects, the Sun cyclically influences the terrestrial ionospheric structure and interplanetary structure. A number of empirical or semi-empirical methods for estimating solar cycle progression exist. The earliest method used relies upon the sunspot number, following the discovery by Wolf (1852) of the 11 yr periodicity in sunspot activity. The "geomagnetic precursor" methods, relying upon measurements of changes in the Earth's magnetic field, determine correlations between sunspot number at solar maximum and the geomagnetic *aa* index at the preceding minimum (e.g., Ohl & Ohl 1979; Feynman 1982; Thompson 1993). Additionally, the solar radio emission at 10.7 cm (F10.7) is a consistent measurement that has been recorded daily since 1947 and is also found to follow the solar activity cycle (Feynman 1982).

Combinations of these techniques have been used to predict the intensity and date of the solar maximum of the current solar cycle. The solar cycle 24 prediction panel[3] (Biesecker & Prediction Panel 2007), led by NOAA, examined several techniques and predicted a maximum in 2013 May that would be weak compared to recent solar cycles. Similarly, recent work presented in Pesnell (2014) predicts a solar cycle 24 maximum F10.7 of no stronger than average and likely weaker than recent solar cycles.

In this Letter, we present a novel approach for determining the solar cycle peak and duration. The solar X-ray background, like other tracers such as the sunspot number and solar radio emission, rises during active times and declines in quiet times. Through an analysis of the X-ray data from the past few solar cycles, we predict the maximum X-ray background level, date of solar maximum, and length of solar cycle 24. In Section 2, we describe our analysis. In Section 3, we compare the X-ray background results to the monthly sunspot number. Section 4 includes discussion of our results.

2. DETERMINATION OF THE SOLAR CYCLE MAXIMUM THROUGH THE X-RAY BACKGROUND

To make our solar cycle predictions, we analyzed *Geostationary Operational Environmental Satellites* (*GOES*) X-ray observations obtained from NOAA's NGDC.[4] We determined the 1–8 Å (corresponding to ~1.5–12.4 keV) background levels using 1 minute data from 1986 through 2014 May 15. The data were obtained from GOES-6, -7, and -8 (solar cycle 22), GOES-8 and -10 (solar cycle 23), and GOES-14 and -15 (2009–present).

The X-ray background was computed as the smoothed minimum flux in a 24 hr time period preceding each 1 minute *GOES* observation. In detail, we use the technique of Hock et al. (2013), which includes the following steps: (1) compute the hourly median with a sliding 1 hr window, (2) determine the instantaneous background as the minimum of these hour medians in the previous 24 hr, and (3) smooth the instantaneous background by the previous 2 hr. The background was computed for both the 1–4 Å and 1–8 Å *GOES* observations. The harder X-ray emission shows no discernible solar cycle trends when compared with the soft X-ray emission, which is the focus of this Letter.

In order to determine the solar maximum and the length of the solar cycle for cycles 22–24, we fit a simple Gaussian to the X-ray background of each solar cycle. We chose a Gaussian for its simplicity in requiring only three free parameters and for its ability to reproduce the shape of the data over a solar cycle. To fit the data, we converted the date and time of the observation into decimal years from the start of the solar cycle (*SCY*). We identified solar cycle 22 as beginning in 1986 August and ending by 1996 May; solar cycle 23 as beginning in 1996 May and ending by 2008 December; and solar cycle 24 as beginning in 2008 December. We then fit a Gaussian of the form:

$$F(\text{SCY}) = F_0 \exp((\text{SCY} - \text{Solar Max})^2/(2\sigma^2)), \quad (1)$$

to the X-ray background. In the equation, F is the logarithm of the X-ray background flux in W m^{-2}, F_0 is the logarithm of

---

Table 1
Results from Gaussian Fits to the 1–8 Å X-Ray Background from *GOES* Using a Variety of Binning Widths

Solar Cycle	F_0 (log W m^{-2})	Solar Max (SCY)	σ (yr)	End Cycle	χ^2
One month					
22	−5.95	1990 Dec (4.27)	6.35	1997 Apr	22.2/94
23	−6.16	2001 Sep (5.38)	7.69	2009 Jun	24.4/118
24	−6.28	2014 Jun (5.53)	7.75	2022 Mar	7.4/35
Two weeks					
22	−5.96	1990 Dec (4.27)	6.39	1997 Apr	82.7/189
23	−6.19	2001 Sep (5.36)	8.05	2009 Oct	99.2/239
24	−6.29	2014 Jan (5.11)	6.94	2020 Dec	24.0/72
One week					
22	−5.97	1990 Dec 12 (4.28)	6.40	1997 May	326.5/387
23	−6.21	2001 Sep 16 (5.38)	8.18	2009 Nov	386/478
24	−6.31	2013 Nov (5.00)	6.76	2020 Sep	107.0/153
0.5 weeks					
22	−5.98	1990 Dec 17 (4.29)	6.46	1997 Jun	1251.2/774
23	−6.21	2001 Sep 13 (5.37)	8.24	2009 Dec	1499.3/956
24	−6.32	2013 Oct (4.82)	6.25	2020 Jan	Oversampled
One day					
22	−5.97	1990 Dec 17 (4.29)	6.46	1997Jun	Oversampled
23	−6.21	2001 Sep 2 (5.34)	8.29	2009 Dec	Oversampled
24	−6.30	2013 Jun (4.51)	5.17	2018 Aug	Oversampled

Notes. The solar cycle, peak flux, solar maximum date, and the corresponding decimal years since the beginning of the solar cycle (SCY), half-width (σ) of the solar cycle, date of the end of the cycle, and χ^2 from the model fit are given. Cases where reduced $\chi^2 \gg 1$ are indicated as "oversampled."

the X-ray background flux at solar maximum in W m^{-2}, *SCY* is the solar cycle year in years, *Solar Max* is the fitted solar maximum value in years from the start of the solar cycle, and σ is the half-width of the solar cycle. In the fitting process, we filtered out any data points with background levels below 10^{-9} W m^{-2}. Such measurements are below the *GOES* 1–8 Å threshold of 3.7×10^{-9} W m^{-2}. The Levenberg–Marquardt algorithm (Levenberg 1944; Marquardt 1963) was used to find the best-fit parameters F_0, *Solar Max*, and σ with the SciPy optimization library in Python.

We determined the effect of the choice of bin size on the solar cycle parameters by computing best-fit values and χ^2 statistics for bin sizes of one month, two weeks, one week, 0.5 weeks, and one day. The best-fit parameters for each solar cycle examined and for each binning level are listed in Table 1. We find that the peak background flux is the most stable parameter, with very little variation in this parameter regardless of bin size. For solar cycles 22 and 23, the solar maximum calculation is also stable, but the duration of the cycle varies by two months for cycle 22 and six months for cycle 23. We tested goodness of fit with the χ^2 statistic, defined as $\chi^2 = \Sigma(\text{observed} - \text{model})^2/\text{std}^2$, where std is the standard deviation of the measurements. The ideal case is where the reduced χ^2 statistic, χ^2 divided by the degrees of freedom (the number of data points fitted minus the number of free parameters fit by the model), is closest to one. For the smaller bin sizes, cases where the reduced χ^2 value is much greater than one are labeled as oversampled. For the largest bin sizes, the standard deviation is large, causing reduced $\chi^2 \ll 1$. The optimized reduced χ^2 values in Table 1 correspond to the one-week binning. The best-fit parameters from the one-week bin size are shown in Table 2. The median one-week background and best-fit Gaussians are shown for solar cycles 22–24 in Figure 1.

Traditional measures of the solar cycle such as sunspot numbers show a double peak due to the solar activity in the northern and southern hemispheres (e.g., Roy 1977). Similarly, the X-ray observations also show the double-peak profile. However, our choice of binning size affects whether the double-peaked structure is blurred or distinct. For this reason, we chose to fit only a single Gaussian to derive the solar maximum and duration, but determined the peaks from examination of the one-week binned data. In Table 2, peak 1 corresponds to the peak in the X-ray background occurring before the fitted solar maximum and peak 2 is the peak following the solar maximum.

Since the current solar cycle 24 is incomplete, the resulting fewer measurements lead to more variability in the fitted solar maximum and duration parameters depending on the chosen bin size. In all cases (Table 2), however, we find that we have reached or passed the solar maximum. Solar cycle 24 is likely to end around 2020, with a maximum uncertainty of 2 yr.

3. COMPARISON TO THE SUNSPOT NUMBER

The earliest method used to determine the solar activity level was through observing changes in the sunspot number. To determine how the X-ray background compares to sunspot number, we analyzed data of the monthly sunspot number from the Solar Influences Data Analysis Center in Belgium.[5] In Figure 2, we show the sunspot number on the same scale as the one-month averaged 1–8 Å X-ray background for solar cycles 22–24. We fit Gaussians to the sunspot number data for each of the solar cycles 22–24, with the best-fit parameters listed in Table 3.

---

[5] The monthly sunspot number was obtained from NASA Marshall Space Flight Center's compilation available here: http://solarscience.msfc.nasa.gov/greenwch/spot_num.txt.
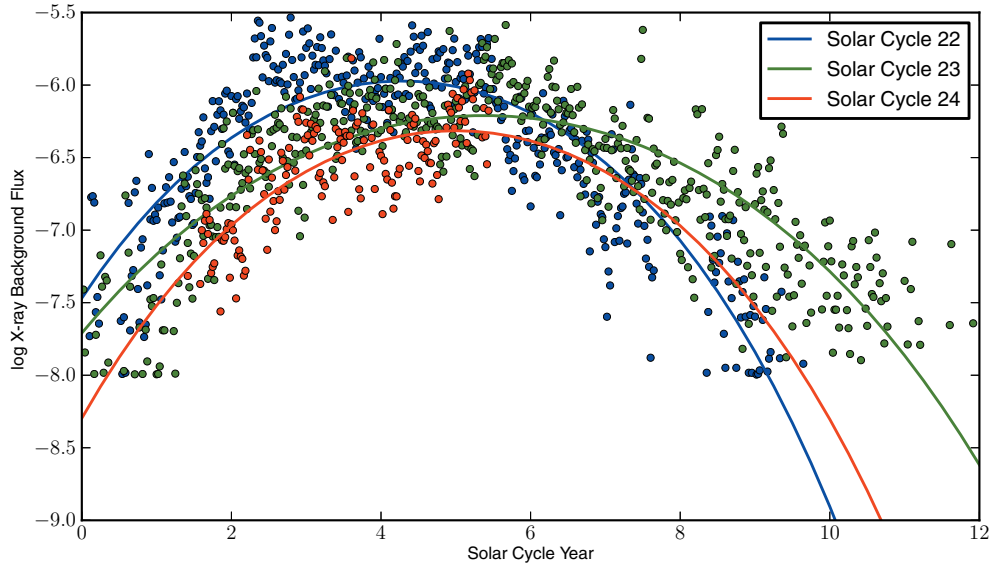
Figure 1. Results of a Gaussian fit (lines) to the one-week averaged 1–8 Å X-ray background (points; in units of log of W m^{-2}) for solar cycles 22–24. The X-ray background flux varies within the solar cycle, with higher values by a factor of 100 from solar minimum to solar maximum.

(A color version of this figure is available in the online journal.)

Table 2
Results from Gaussian Fits to the 1–8 Å X-Ray Background from *GOES* with the One-week Binning

Solar Cycle	F_0 (log W m^{-2})	Solar Max (SCY)	Peak 1 (log W m^{-2})	Peak 1 Date	Peak 2 (log W m^{-2})	Peak 2 Date
22	−5.97	1990 Dec (4.28)	−5.59	1989 Jun	−5.70	1991 Apr
23	−6.21	2001 Sep (5.38)	−5.72	2000 Jul	−5.87	2002 Jan
24	−6.31	2013 Nov (5.00)	−5.82	2012 Jul	−5.92	2014 Feb

Note. The solar cycle, hemisphere-averaged solar maximum flux, hemisphere-averaged solar maximum date, and corresponding value in solar cycle years (SCY) or decimal years since the beginning of the solar cycle, and the flux and date of each of the two peaks in each cycle are given.

Table 3
Details of the Monthly Sunspot Number

Solar Cycle	SS Max No.	SS Max Date	Peak 1 No.	Peak 1 Date	Peak 2 No.	Peak 2 Date
22	162	1990 Jul	196	1989 Jun	173	1991 Jul
23	120	2001 Mar	170	2000 Jul	150	2001 Sep
24	75	2013 Nov	96	2011 Nov	102	2014 Feb

Notes. The hemisphere-averaged sunspot (SS) Max values correspond to best-fit median date from a Gaussian fit and the sunspot number at that date. Peak 1 corresponds to the peak occurring before the fitted median and Peak 2 corresponds to the peak level after the fitted median (SS Max).

In addition to the best-fit solar maximum date and associated sunspot number, Table 3 also includes the date and sunspot number of each of the peaks, obtained from analysis of the sunspot number data. Peak 1 is the peak sunspot number before solar maximum and peak 2 is the peak sunspot number following solar maximum. Given the difference in the binning method for the sunspot number from the X-ray background, we expect some variability in comparing the dates of solar max. We do find reasonable agreement, however, even when comparing with the one-week binning of the 1–8 Å X-ray background (Table 2). For cycles 22 and 23, the X-ray background solar maximum (regardless of binning) is within about four months from the sunspot number solar maximum. The cycle 24 estimates agree for peak 2 (2014 February; which is the peak to date of the

analysis), which was the maximum in sunspot number to date, although there is a large difference in the peak 1 dates of eight months.

The peaks in the sunspot number and X-ray background are both higher preceding the solar max for cycles 22 and 23. The sunspot peak numbers in cycle 22, 192 for peak 1, and 173 for peak 2, were higher than in subsequent cycles. The overall highest peaks in sunspot number from cycles 23 and 24 are 13% and 48% lower than the peak in cycle 22, respectively.

4. DISCUSSION

The solar soft X-ray emission is an important indicator of the state of the corona. While the mechanisms of coronal heating are poorly understood, the process is connected with solar magnetic activity (e.g., Vaiana & Rosner 1978). Previous soft X-ray studies have shown that variations exist in the derived luminosity from minimum to maximum, by a ratio of five to six times (Judge et al. 2003). With uniform observations over the past nearly three solar cycles, the *GOES* soft X-ray measurements provide a powerful database for characterizing the coronal variability and a tool for not only monitoring of flare activity but also for space weather forecasting.

Based on our analysis of the *GOES* 1–8 Å (∼1.5–12 keV) observations from 1986 to the present, we have confirmed that the X-ray emission varies with solar cycle. We determined a soft X-ray background as the minimum flux in a 24 hr time
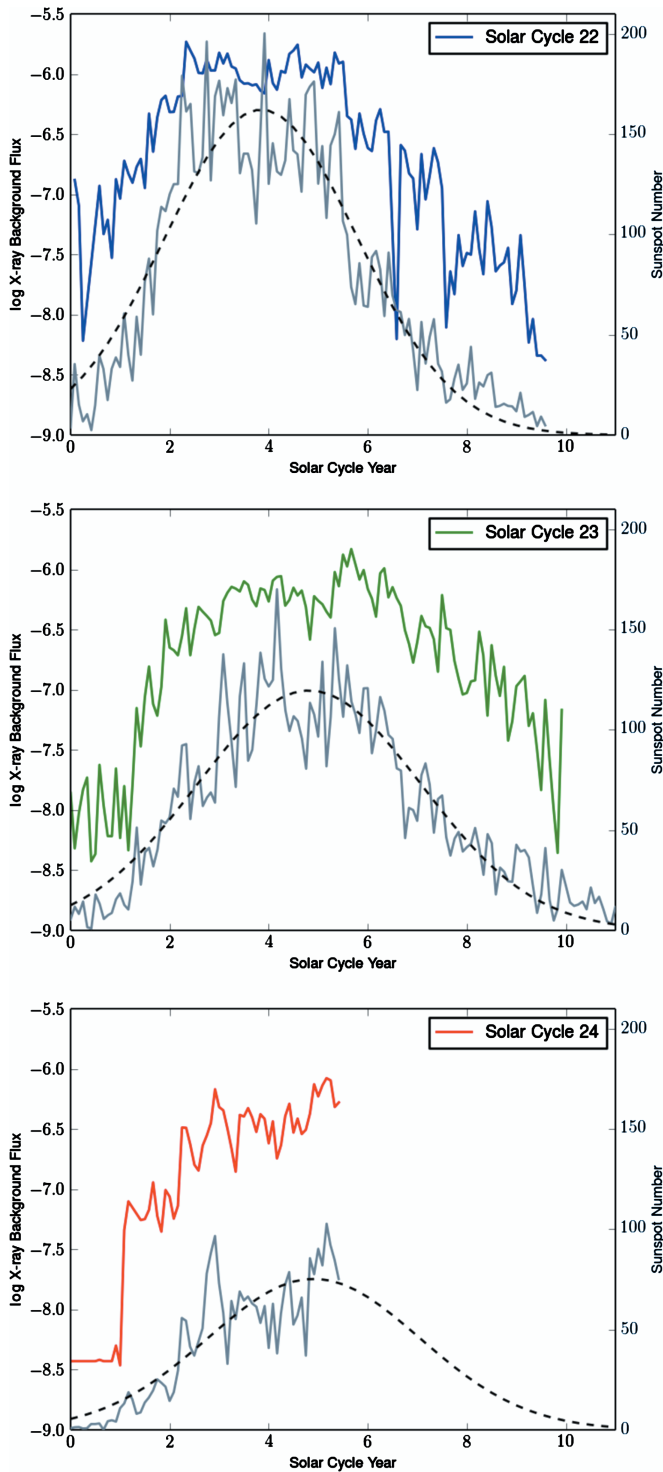
Figure 2. Comparison of the one-month averaged 1–8 Å X-ray background (colored lines; in units of log of $W\,m^{-2}$) to the monthly sunspot number (gray line) for solar cycles 22–24. The dashed lines show the Gaussian fit to the sunspot number. Note that color notations are the same as in Figure 1.

(A color version of this figure is available in the online journal.)

period preceding each 1 minute *GOES* observation. From our analysis, we show that the variance in this X-ray background follows a cyclical pattern from solar minimum to maximum. Additional variations between solar cycles (e.g., differences in the solar maximum flux and length of the cycle) are also found, with the peak background at solar maximum declining over the

past two cycles. In particular, we find that the solar cycle 22 X-ray background peak of $1.1 \times 10^{-6}\,W\,m^{-2}$ is 1.6 times the solar cycle 23 peak. The predicted peak for solar cycle 24 is $5.21 \times 10^{-7}\,W\,m^{-2}$, 25% lower than the peak background level in cycle 23 and only half of the peak level in cycle 22. This variance is consistent with the variability found in the sunspot cycle during the same time periods, as shown in Section 3.

Further, we find that the soft X-ray emission during solar minimum has also declined over the past two cycles. The average level over the year of solar minimum preceding each solar cycle declined by a factor of ~4 from $8.3 \times 10^{-8}\,W\,m^{-2}$ during solar cycle 22 to $2 \times 10^{-8}\,W\,m^{-2}$ during solar cycle 23. During solar cycle 24, the solar minimum average is unable to be determined reliably, since 72% of the measurements in solar minimum were below the *GOES* threshold of $3.72 \times 10^{-8}\,W\,m^{-2}$. However, this evidence shows that the background was lower than the previous minimum, consistent with results from more sensitive soft X-ray instruments such as the SphinX X-ray spectrophotometer on the Russian CORONAS-PHOTON spacecraft (e.g., Sylwester et al. 2012).

These results carry merit as a historical study of the solar X-ray emission. Additionally, they exhibit the potential of our technique for space weather climatology. One use is as an alternative method in determining the characteristics of the solar cycle. Based upon our results, we predict the hemisphere-averaged maximum for solar cycle 24 as occurring in 2013 November, with the peak so far having occurred in 2014 February. The X-ray-based predictions we made for the previous two solar cycles, along with the current cycle, were in good agreement with the sunspot cycle. Our analysis also allows us to estimate the end date of the current cycle. Our predicted end dates for solar cycles 22 and 23 were 1997 April and 2009 November, respectively. In cycle 23, our predicted end date is ~1 yr later than NOAA SWPC's agreed upon date of 2008 December. Our predicted end date for solar cycle 24 is 2020 September.

Additionally, since the occurrence of X-ray flares is linked to solar activity and we have shown that the soft X-ray background scales with this, the X-ray background may also prove an important tool in X-ray flare forecasting. The intensity and number of flares, for instance, is also shown to scale with solar cycle. In future work, we will explore the use of this technique as a diagnostic for in-progress flare forecasting. We also plan to compare these solar cycle measures to those from observations with the coronagraph at the John W. Evans Solar Facility of the National Solar Observatory at Sacramento Peak (Altrock 2014).

REFERENCES

Altrock, R. C. 2014, SoPh, 289, 623
Biesecker, D. A., & Prediction Panel, S. C. 2007, BAAS, 39, 210
Feynman, J. 1982, JGR, 87, 6153
Hock, R. A., Woodraska, D., & Woods, T. N. 2013, SpWea, 11, 262
Judge, P. G., Solomon, S. C., & Ayres, T. R. 2003, ApJ, 593, 534
Levenberg, K. 1944, QApMa, 2, 164
Marquardt, D. 1963, SJAM, 11, 431
Ohl, A. I., & Ohl, G. I. 1979, in NOAA Solar–Terrestrial Predictions Proceedings, A New Method of Very Long-term Prediction of Solar Activity, Vol. 2., ed. R. F. Donnelly (Boulder, CO: NOAA), 258
Pesnell, W. D. 2014, SoPh, 289, 2317
Roy, J.-R. 1977, SoPh, 52, 53
Sylwester, J., Kowalinski, M., Gburek, S., et al. 2012, ApJ, 751, 111
Thompson, R. J. 1993, SoPh, 148, 383
Vaiana, G. S., & Rosner, R. 1978, ARA&A, 16, 393
Wolf, M. 1852, MNRAS, 13, 29

DISTRIBUTION LIST

DTIC/OCP
8725 John J. Kingman Rd, Suite 0944
Ft Belvoir, VA 22060-6218   1 cy

AFRL/RVIL
Kirtland AFB, NM 87117-5776   2 cys

Official Record Copy
AFRL/RVBXS/Dr. K. S..Balasubramaniam 1 cy

This page is intentionally left blank.